



Disturbance Feedforward Control for Vibration Suppression in Adaptive Optics of Large Telescopes

Martin Glück, Jörg-Uwe Pott, Oliver Sawodny



Institute for System Dynamics



Disturbances affecting the Telescope Resolution

- Atmospheric Turbulences
- Structural Vibrations



Adaptive Optics

Observations with Faint Guide Stars

- Long exposure for better Signal-to-Noise-Ratio
- Adaptive Optics (AO) loop speed is reduced
- High frequency Vibrations are seen at the Telescope Mirror (> 5 Hz)
- Classical AO loop is to slow



Accelerometer-based Disturbance Feedforward





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CAccelerometer-based Disturbance Feedforward Control

Compensation of the Telescope Vibrations

- Measuring vibrations at relevant telescope mirrors with additional accelerometers
- Reconstruction of the Optical Modes influenced by Vibrations (Piston, Tip, Tilt Zernike Modes)
- Compensating by the Tip-Tilt mirror and Deformable Mirror
 - Independent of the Wavefront Exposure Time
 - Suppression of high frequency Vibrations











1. Methods for Vibration Suppression

- 2. Simulation Results
- 3. Conclusion and Outlook











Assumptions

- > DM \approx 1 (fast position control)
- 2 sample delay of the WFS

Disturbance Modelling

- > Atmosphere
 - > Approximation of the temporal power spectral density by a second order AR model

$$X_t = a_1 X_{t-1} + a_2 X_{t-2} + \epsilon_t$$

- Vibrations
 - Modal representation of the mechanical system (considering dominant natural frequencies) $\ddot{x} + 2d\omega_0\dot{x} + \omega_0^2x = u$



Concept

- Measuring Vibrations with Accelerometers at Telescope structure
- Reconstructing Optical Modes (Piston, Tip, Tilt, Defocus)
 - Different Tip-Tilt Sensitivity of each Mirror
- Adding to the Control Input of the Adaptive Mirrors (M4, M5)
 - Considering time delay





Challenges

- Online Double Integration of Accelerometer Signals
 Unstable
- ➢ Time Delay

Reconstruction Methods

- Approximation of a Double Integrator by a Bandpass Filter, Böhm[2]
- Disturbance Observer
 - Modal Model of the Mechanics
 - Luenberger Observer Design
- Adaptive Resonator , Keck[1]
 - Using Online Fourier Analysis

$$\ddot{\varphi}_{AR,i} = A \cos(\omega_1 t) + B \sin(\omega_1 t)$$
$$\varphi_{AR,i} = -\frac{1}{\omega_1^2} \ddot{\varphi}_{AR,i}$$



$$\dot{x} = \begin{bmatrix} 0 & 1 \\ -\omega_0^2 & 0 \end{bmatrix} x + \begin{bmatrix} 0 \\ 1 \end{bmatrix} u \quad y = \begin{bmatrix} -\omega_0^2 & 0 \end{bmatrix} x$$
$$\dot{\hat{x}} = A\hat{x} + L (y - \hat{y}), \qquad \hat{y} = C\hat{x}$$

$$A = \int_0^t g(\ddot{z}_i - \ddot{z}_{AR,i}) \cos(\omega_1 \tau) d\tau$$
$$B = \int_0^t g(\ddot{z}_i - \ddot{z}_{AR,i}) \sin(\omega_1 \tau) d\tau$$









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Simulation Parameters

- Wind velocity 10 m/s
- Natural frequenies at 10 Hz and 13Hz
- Sample Rate Accelerometers 1 kHz
- Disturbance Feedforward Control with a Kalman Reconstructor







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Bright guide star 10 mag und exposure time 800 Hz

- > Typical atmospheric condition 0.8 arcsec
- > Exciting the system by varying the natural frequency 0 Hz ... 50 Hz





Faint guide star with 14.6 mag and exposure rate 200 Hz

> Exciting the system by varying the natural frequency 0 Hz ... 50 Hz





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Conclusion

- High frequency vibrations worsen the performance for observations with faint guide stars
- Improving the performance with an Accelerometerbased Disturbance Feedforward control

Outlook

- Considering the influence of the actuator dynamics (M4, M5)
- Testing Disturbance Feedforward at the LBT



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Thank You! Questions ?



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